

# The magnetosphere under weak solar wind forcing

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**Abstract.** The Earth's magnetosphere was very strongly disturbed during the passage of the strong shock and the following interacting ejecta on 21–25 October 2001. These disturbances included two intense storms ( $D_{st}^* \approx -250$  and  $-180$  nT, respectively). The cessation of this activity at the start of 24 October ushered in a peculiar state of the magnetosphere which lasted for about 28 h and which we discuss in this paper. The interplanetary field was dominated by the sunward component [ $\mathbf{B}=(4.29 \pm 0.77, -0.30 \pm 0.71, 0.49 \pm 0.45)$  nT]. We analyze global indicators of geomagnetic disturbances, polar cap precipitation, ground magnetometer records, and ionospheric convection as obtained from SuperDARN radars. The state of the magnetosphere is characterized by the following features: (i) generally weak and patchy (in time) low-latitude dayside reconnection or reconnection poleward of the cusps; (ii) absence of substorms; (iii) a monotonic recovery from the previous storm activity ( $D_{st}$  corrected for magnetopause currents decreasing from  $\sim -65$  to  $\sim -35$  nT), giving an unforced decreased of  $\sim 1.1$  nT/h; (iv) the probable absence of viscous-type interaction originating from the Kelvin-Helmholtz (KH) instability; (v) a cross-polar cap potential of just 20–30 kV; (vi) a persistent, polar cap region containing (vii) very weak, and sometimes absent, electron precipitation and no systematic inter-hemisphere asymmetry. Whereas we therefore infer the presence of a moderate amount of open flux, the convection is generally weak and patchy, which we ascribe to the lack of solar wind driver. This magnetospheric state approaches that predicted by Cowley and Lockwood (1992) but has never yet been observed.

**Keywords.** Magnetospheric physics (Magnetospheric configuration and dynamics; Solar wind-magnetosphere interactions; General or miscellaneous)

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## 1 Introduction

The terrestrial magnetosphere interacts continuously with the magnetized solar wind flow. The basic form of interaction is a compression/dilation caused by a variable solar wind dynamic pressure. Major forms of geomagnetic activity depend crucially on the orientation of the interplanetary magnetic field (IMF). For southward IMF, these include enhanced plasma convection (Dungey, 1961; Cowley, 1982); erosion of magnetic flux from the dayside and its transport to the nightside, producing a dilated polar cap (Burch, 1972); episodes of explosive energy conversion in the geomagnetic tail called substorms (McPherron et al., 1973); and an energized ring current during storms. For northward IMF, reconnection at high latitudes (Dungey, 1963; Russell, 1972; Maezawa, 1976) may lead to cross-polar cap potentials of several tens of kV (Cumnoek et al., 1992; Freeman et al., 1993; Knipp et al., 1993) while viscous-type interactions (Axford and Hines, 1961), most notably the Kelvin-Helmholtz instability (Dungey, 1955; Southwood 1979), can excite geomagnetic pulsations (Southwood, 1974; Chen and Hasegawa, 1974). Is the magnetosphere ever in a “ground” state where this diverse activity is switched off or operating at a very low level? In this paper we report on a time interval on 24–25 October 2001 when the IMF was pointing very nearly sunward with smaller and variable  $B_y$  and  $B_z$  components such that little solar wind-magnetosphere coupling might be expected. In fact, the magnetosphere appeared to be variably but near-continuously driven, albeit at a remarkably low rate.

Interestingly, this state was preceded by one of the most disturbed periods in this solar cycle, where two major enhancements of the ring current, each accompanied by a burst of substorm activity, were recorded. The interplanetary cause of this activity was a merger of at least two ejecta driving a strong shock. The passage at Earth of such mergers (often called “complex ejecta” (Burlaga et al., 2001, 2002) are typically of long duration – of the order of several days – and thus

the disturbances they elicit may last for a long time. Here we focus on the ejecta merger which passed Earth on 21–25 October 2001. We shall in this paper characterize the state of the magnetosphere when the intense activity attending the passage of the earlier portion of this configuration, including the shock, had subsided. We shall find almost quiescent conditions with, however, a number of peculiar features which may be a residue of the previous activity.

## 2 Observations

### 2.1 Interplanetary observations

Interplanetary parameters for the 5-day interval 21–25 October 2001, are shown in Fig. 1, which displays plasma and magnetic field measurements made by the ACE spacecraft from near the L1 libration point  $\sim 220 R_E$  (Earth radii) upstream of Earth. The data were acquired by the MAG (magnetic field) and SWEPAM (plasma) instruments at a temporal resolution of 16 s and 64 s, respectively. Plotted are the proton density, bulk speed, temperature, dynamic pressure, the  $\alpha$  particle-to-proton number density ratio, total field strength and its GSM components, and the proton plasma  $\beta$ . The red trace in the third panel is the proton temperature expected for the expanding solar wind resulting from statistical studies (Lopez, 1987).

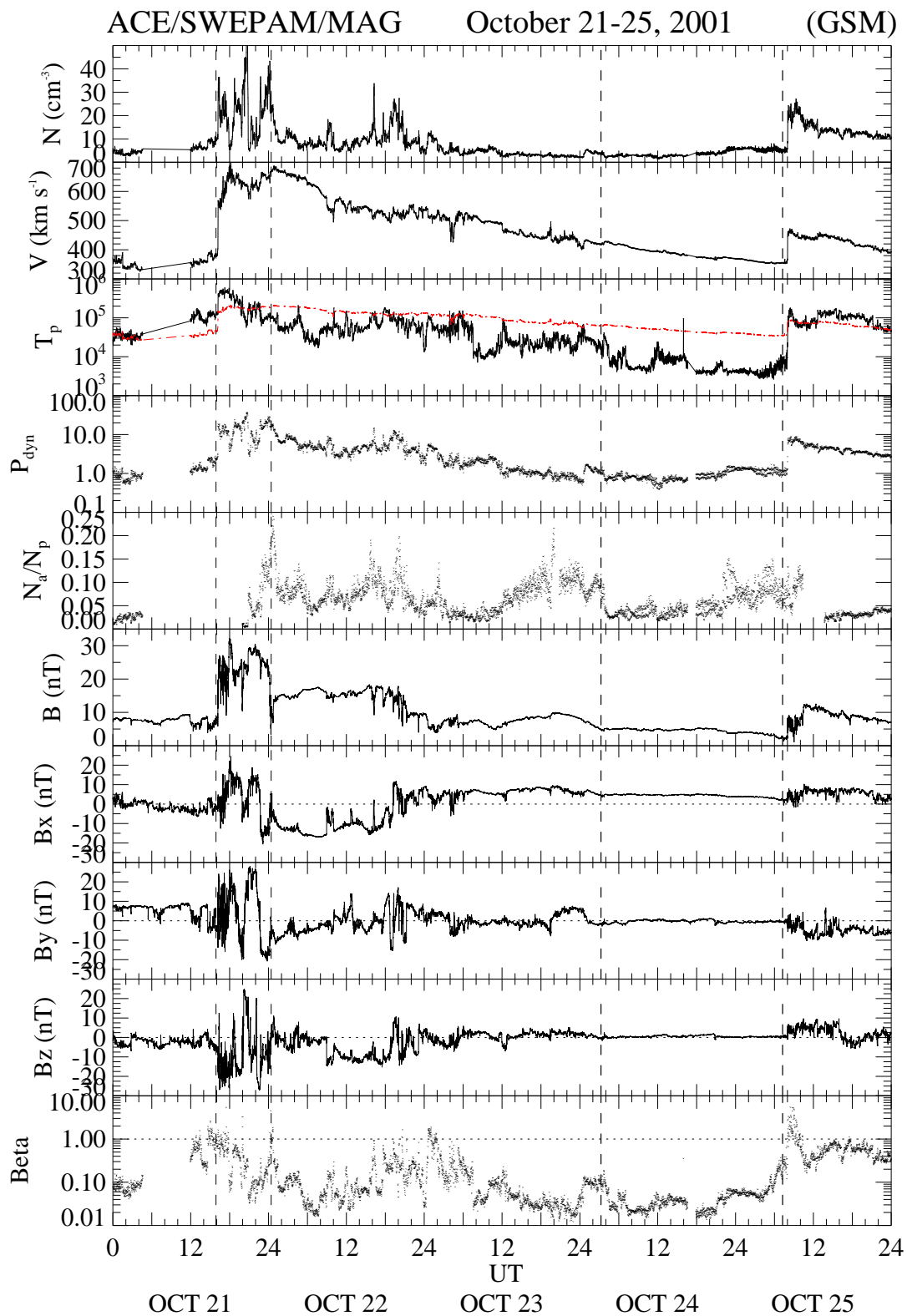
A strong shock (compression ratios in the density and magnetic field strength of  $\sim 3$ ) passes the spacecraft at  $\sim 16:00$  UT on 21 October 2001 (first vertical guideline). Behind the shock very strong magnetic field strengths are observed as is also an undulating density profile reaching very large values ( $\sim 50 \text{ cm}^{-3}$ ). Just after the last of these, the ejecta is entered (second vertical guideline). The time of entry may be inferred from the enhancement in the alpha/proton number density ratio, the generally strong fields, and the low proton temperatures (where by “low” we understand lower than the expected temperature), and proton  $\beta < 1$ . These features continue – with a notable exception during  $\sim 01:00$ – $02:00$  UT, 23 October, when there is a magnetic field strength decrease and  $\beta$  rises above unity – up to the arrival of a second shock at  $08:00$  UT, 25 October. One may note the generally declining  $V$ -profile (panel 2) from after the arrival of leading shock to the arrival of a second shock advancing into the ejecta merger. This simple speed profile is one of the features characterizing complex ejecta (see example 1 in Burlaga et al., 2002). The magnetic field in the interval of interest ( $03:25$  UT, 24 October– $07:25$  UT, 25 October at ACE; bracketed by the last two vertical guidelines) is remarkable in that it is dominated by a sunward ( $B_x$ ) component with smaller and variable east-west and north-south components ( $B_x = 4.29 \pm 0.77$  nT;  $B_y = -0.30 \pm 0.71$  nT, and  $B_z = 0.49 \pm 0.45$  nT). Further striking features are a field of low fluctuation ( $B = 4.40 \pm 0.78$  nT) in a plasma of very low proton temperature ( $T = 6600 \pm 2710$  K) and  $\beta$  ( $\sim 0.04$ ) con-

taining periods of elevated  $\alpha$ -particle concentrations with respect to the protons (up to 15%). The dynamic pressure ( $= 0.94 \pm 0.25$  nPa) is substantially lower than the long-term average of 2.0 nPa, and fairly steady. These quiet values may be contrasted with the earlier interval on 21–23 October, where high and fluctuating densities, strong magnetic fields, and sporadically high temperatures may be seen. Studying the coalescence, total or partial, of ejecta in the inner heliosphere leading to complex ejecta, Farrugia and Berdichevsky (2004) showed that ejecta mergers may strengthen the leading shock, compress the field and plasma of the leading ejecta, and heat the plasma. These are all factors which tend to intensify the geoeffectiveness of interplanetary configurations. This we believe is the cause of the contrast between the earlier and later intervals. The interval of increasingly quiescent geomagnetic activity ends abruptly at  $08:00$  UT, 25 October, when, as noted, a shock advances past ACE.

### 2.2 Global indicators of disturbance

Major indices of geomagnetic activity are shown in Fig. 2. From top to bottom the panels display the  $D_{st}$  storm index corrected for the effect of magnetopause currents ( $D_{st}^*$ ), the 3-hourly  $K_p$  index on a scale from 0 to 8; the auroral electrojet indices AE and AL; and two measures of the cross-polar cap potential, one derived from SuperDARN convection measurements (see Sect. 2.5) (black curve) and another from a statistical analysis of DMSP data reported by Boyle et al. (1997; their Eq. 1). The AL index characterizes auroral activity, a sharp drop by several hundred nT usually denoting the presence of ionospheric currents associated with substorms in the geomagnetic tail. Vertical guidelines delimit our study interval, where an estimated 45 min propagation delay time from ACE to the ground has been taken into account. In the earlier interval there occur two strong enhancements of the ring current indicated by the two minima in  $D_{st}^*$ , two bursts of substorm activity during the peaks in  $D_{st}^*$  and  $K_p$ , and a cross-polar cap potential maximizing at  $\sim 120$  kV (black trace). (The results of Boyle et al. are not expected to agree with measured values at high solar wind electric fields, as in this early interval, because saturation sets in Russell et al., 2001; Siscoe et al., 2002, at values typically above  $\sim 6 \text{ mV m}^{-1}$ ; Mühlbacher et al., 2005.) By contrast, in the interval of interest here, aside from a quasi-linear recovery of the  $D_{st}^*$  from the very strong, earlier activity, the other indices show remarkably steady and low level activity. Particularly noteworthy are the low values of the cross-polar cap potential, ranging from below 20 kV to  $\sim 30$  kV. By measuring the rate of plasma convection within the magnetosphere (Reiff et al., 1981; Stern, 1984), this quantity is a good indicator of the strength of solar wind coupling. We now seek to characterize this state of the magnetosphere and the conditions under which it occurred.

The substorm indices above are average quantities. To check directly whether there is a return of open flux to the



**Fig. 1.** Interplanetary plasma and magnetic field observations from the ACE spacecraft for the 5-day period 21–25 October 2001. The panels show, from top to bottom, the proton density, bulk speed, temperature, and dynamic pressure, the alpha particle-to-proton number density ratio, the total field strength, the components of the magnetic field vector in GSM coordinates, and the proton plasma beta.

dayside leading to a systematic shrinkage of the polar cap during our interval, we examined Wind measurements. On 24 October, the spacecraft was crossing the near-Earth tail on the dawnside, moving from  $(-14.0, -5.4, 1.2) R_E$  at 04:00 UT to  $(-14.3, -21.6, 6.7) R_E$  (GSM) at 18:00 UT, at which time it exited into the magnetosheath. The measurements show a magnetic field dominated by its sunward component (i.e. tailward stretched field), with a duskward component  $B_y = 6.0 \pm 1.4$  nT and with  $B_z = 1.7 \pm 1.4$  nT (GSM). The field strength decreased smoothly from  $\sim 24$  nT to  $\sim 7$  nT. There are no signs of abrupt rotations as might be associated with dipolarizations of the magnetic field, i.e. there is no indication of return of open flux at substorm onsets.

### 2.3 Polar cap precipitation

In an effort to see if dayside reconnection is ongoing during our period of interest, we next examine the polar rain. Current notions ascribe this electron precipitation to coronal electrons which propagate in interplanetary space as a closely field-aligned beam (the strahl) and gain access to the polar regions of Earth through reconnection on the dayside (see Fig. 1, Fairfield and Scudder, 1985). The polar cap precipitation has long been known to exhibit a pronounced north-south asymmetry related to the sector structure of the IMF (Yaeger and Frank, 1976). For, as here, a toward sector ( $B_x > 0$ ), this asymmetry should favor the Southern Hemisphere where the reconnected field lines connect back to the Sun.

The Polar spacecraft (orbital period  $\sim 18.5$  h) made four passes over the polar regions, two on each day. Polar was following an orbit lying approximately in the 22:00–08:00 MLT plane, and rising in invariant latitude (ILT) up to  $\sim 88^\circ$ . In the period we study, perigee ( $\approx 1.51 R_E$ ) occurred at  $\sim 13:00$  MLT and apogee ( $\approx 9.55 R_E$ ) at  $\sim 22:00$  MLT. We examine first the passes over the northern polar cap. The relevant data from the HYDRA instrument are plotted in Figs. 3a, b (top and bottom panels, respectively). The panels show the invariant latitude (ILT), the differential energy fluxes (DEF) of downflowing ions and electrons, the integrated energy and number fluxes (electrons in red). In the northern polar cap, downflowing particles move parallel the field, and we have plotted fluxes whose pitch angle lies in the range 0 to  $30^\circ$ . The data, which have been corrected for spacecraft potential effects, have a time resolution of 13.8 s.

The passes over the Northern Hemisphere last many hours, with the spacecraft ascending slowly in ILT to above  $80^\circ$  while descending from apogee. They are in the nightside region (MLT  $\sim 22:00$ – $23:00$ ). On the first pass (Fig. 3a) significant electron DEFs are measured only up to  $\sim 500$  eV, and intensities are very weak starting from 07:00 to 12:00 UT, with intermittent drop-outs ( $\text{DEF} < 10^5 \text{ cm}^{-2} \text{ s}^{-1} \text{ sr}^{-1} (\Delta E/E)^{-1}$ ). Electron number fluxes are less than  $10^6 \text{ cm}^{-2} \text{ sr}^{-1} \text{ s}^{-1}$  and ion number fluxes are less than  $10^5 \text{ cm}^{-2} \text{ sr}^{-1} \text{ s}^{-1}$ , a precipitation form classi-

fied as “void” by Shinohara and Kokubun (1996; their Table 1). The nightside boundary of the polar cap appears to be crossed at  $\sim 07:00$  UT and be located at  $\sim 73^\circ$ .

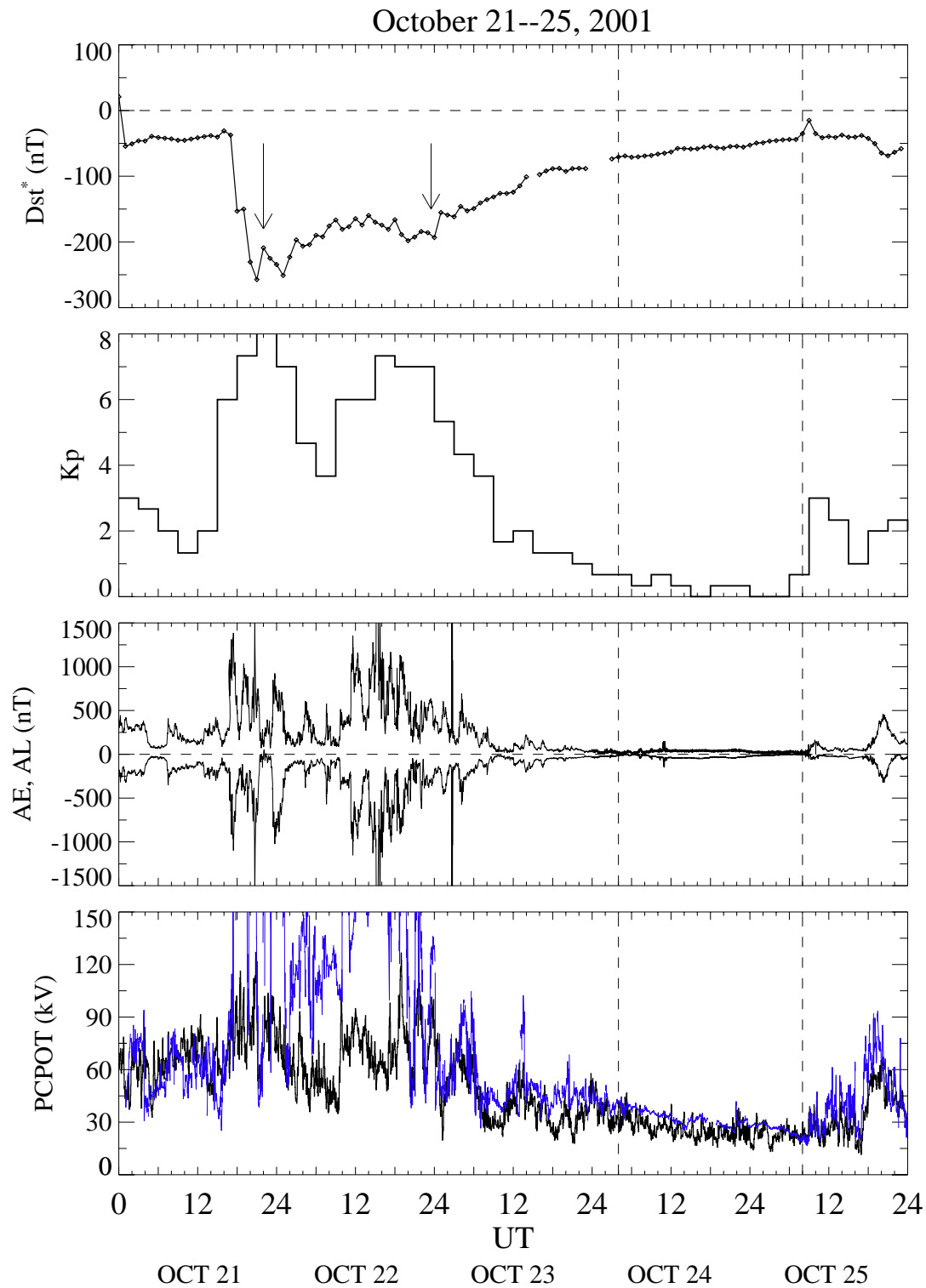
The second pass, from 23:00 UT, 24 October to 07:00 UT, 25 October (Fig. 3b), occurs in a similar MLT-ILT range as the previous pass. Similar precipitation characteristics are seen. The precipitation type may also be classified as void, except during sporadic, moderate enhancements such as occur at  $\sim 01:00$  UT and occasionally in 03:00–04:00 UT. Partial dropouts are also evident, particularly near the highest latitudes later on in the pass. Again, the polar cap boundary at  $\sim 22:30$  MLT appears to be located at  $\sim 73^\circ$  ILT.

We next consider the Southern Hemisphere passes (Figs. 4a, b), plotting this time fluxes antiparallel to the magnetic field ( $150$ – $180^\circ$  pitch angle range). The first pass shows considerably more intense precipitation than in the north, with significant DEFs in energies up to 700 eV. A polar arc (at  $\sim 13:30$  UT) and a burst of enhanced precipitation (at  $\sim 13:48$ – $14:10$  UT) bracket a region at the highest latitudes (above  $-80^\circ$  ILT) where the fluxes are intermittently as weak as in the northern polar cap. It is hard to locate the polar cap boundary. At  $\sim 21:30$  MLT, if we place it at the start of the gradient in electron precipitation at  $\sim 14:48$  UT, it lies at  $\sim -73^\circ$  ILT. At midmorning MLTs, if we place the polar cap boundary at the start of the polar arc observation, it would lie at  $\sim -80^\circ$ .

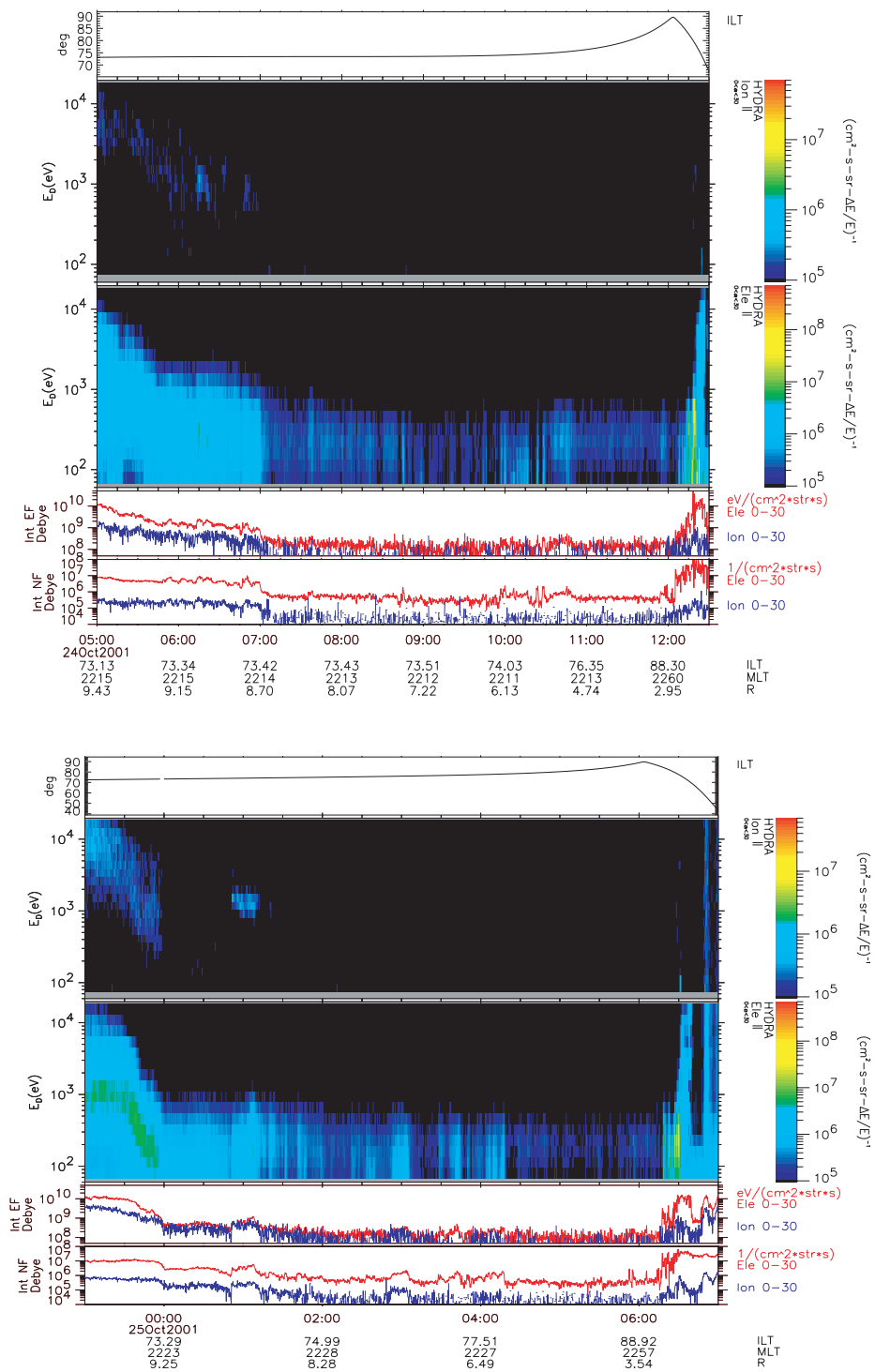
The second southern pass on 25 October shows a marked contrast with the previous one. DEFs extend up to only 500 eV, and the integrated energy fluxes (last-but-one panel) are weaker. There is essentially a void on the nightside portion (to the right). In the interval 07:36–07:45 UT at high ILTs ( $78$ – $88^\circ$ ) on the dayside ( $\sim 08:00$  MLT) intense bursts of precipitation of energies reaching  $\sim 1$  keV are present. These are likely signatures of polar arcs. The polar rain observations are still continuing at a ILT of  $\sim -71^\circ$  when a rapid energization occurs at  $\sim 08:45$  UT, presumably on arrival of the shock seen at ACE  $\sim 45$  min earlier (cf. Fig. 1).

Summarizing, we have a polar cap precipitation which is (i) very weak, and sometimes absent, in the north; (ii) shows some indications of a north-south asymmetry which, however, becomes much attenuated as time progresses; and (iii) is confined to energies of a few hundred eV only. The presence of polar arcs were noted. In general, the observations support the idea of a low level of solar wind-magnetosphere coupling. Two caveats on the interpretation should be borne in mind, both related to the weakness of the precipitation. (1) The lack of obvious north-south asymmetry, in so far as this is a meaningful concept when the precipitation is so weak, may be due to the fact that the ejecta field lines are connected at both ends to the Sun (Makita and Meng, 1987). (2) The extent of the open field line region is hard to pin down. However, an exact knowledge of the extent of the open field line region is not central to the main thrust of this paper.

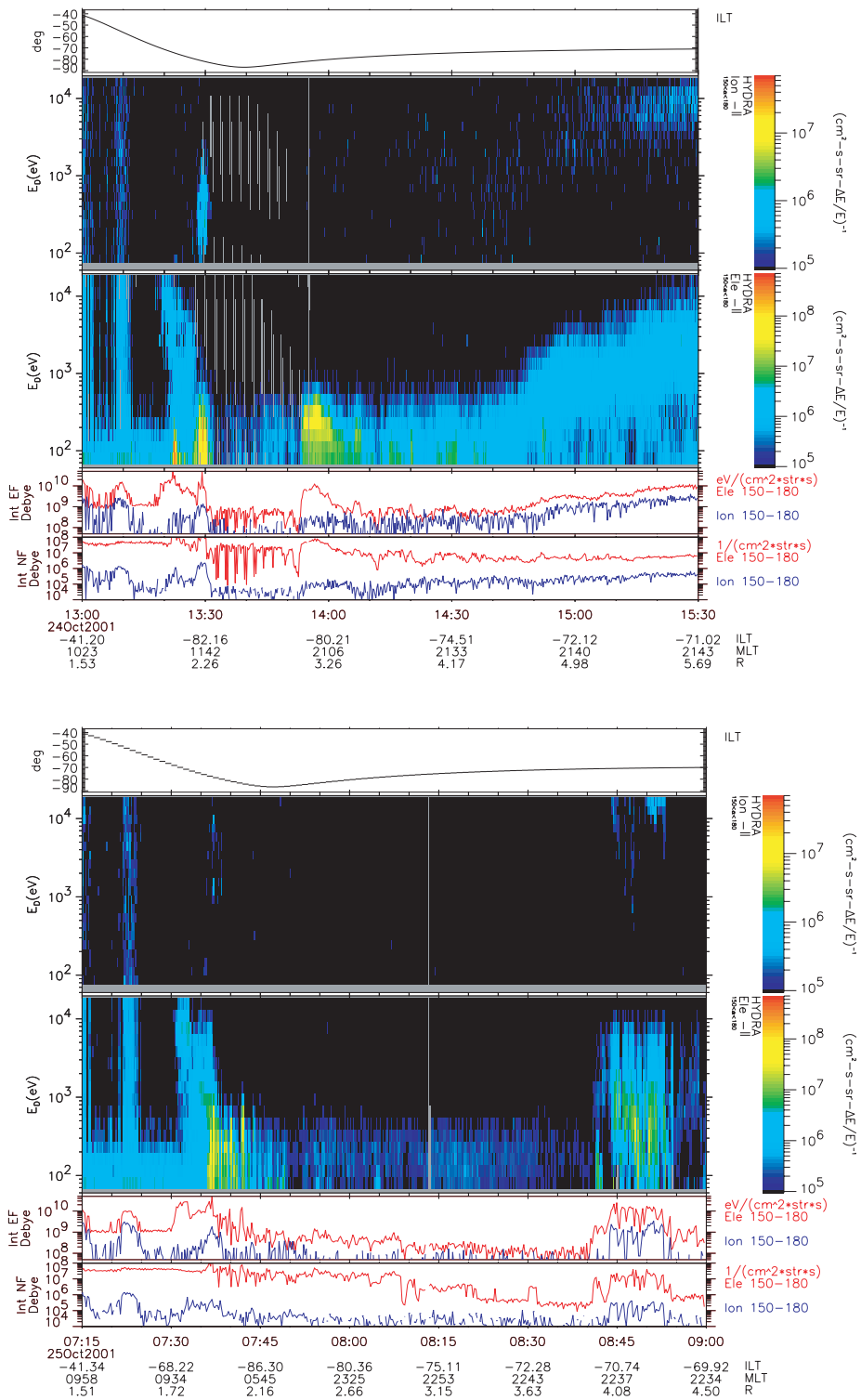
To complement and extend the observations by Polar, we surveyed data acquired at low altitudes by the DMSF



**Fig. 2.** Measures of geomagnetic activity during 21 October–25 October: The  $D_{st}$  index corrected for magnetopause currents (panel 1), the  $K_p$  index (panel 2), the auroral electrojet index AE (positive values) and the AL-index (negative values) (panel 3), and the cross-polar cap potential obtained from SuperDARN convection data (black trace), and that derived from an empirical formula based on DMSP passes according to Boyle et al. (2000) (last panel). The vertical dashed guidelines bracket the interval we study.



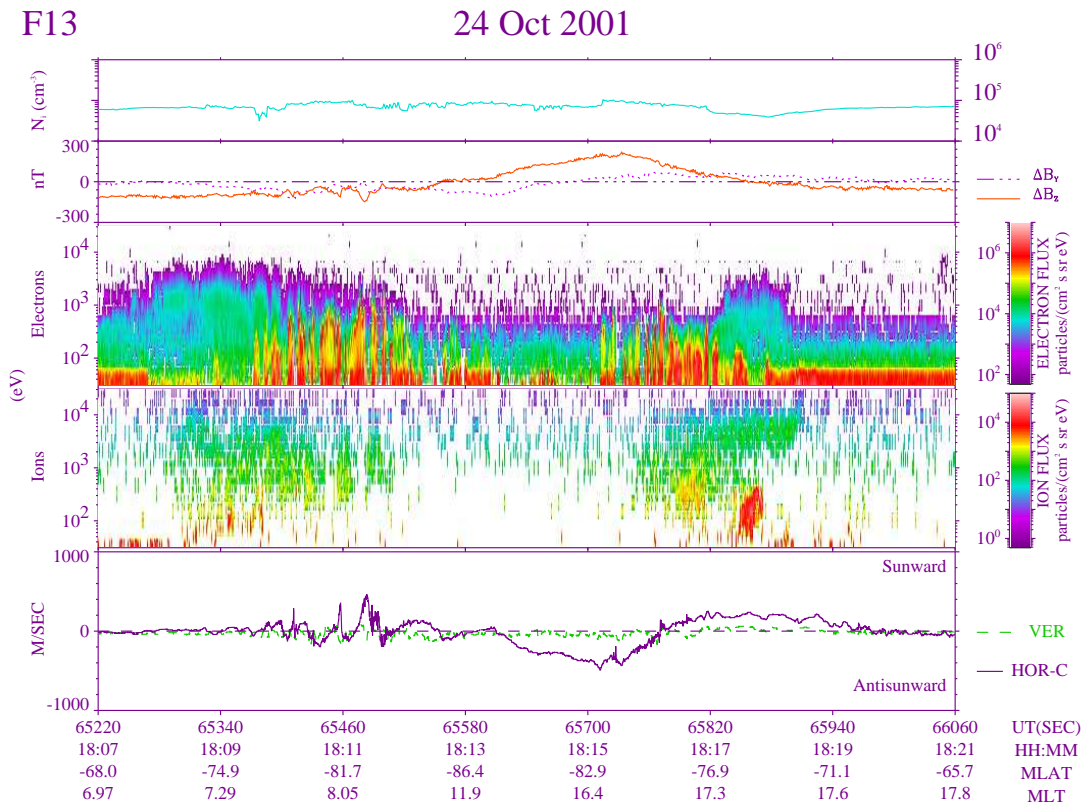
**Fig. 3. (a), (b):** Polar/Hydra observations of particle precipitation at high latitudes. The panels show the invariant latitude of the spacecraft, differential energy fluxes of ions and electrons travelling parallel to the magnetic field, with intensities according to the color bar on the right; integrated energy and number fluxes (electrons in red). Bottom panel: Similar to panel (a) but for the next pass over the Northern Hemisphere.



**Fig. 4. (a), (b):** Similar to Fig. 3, but for the two consecutive passes over the Southern Hemisphere. The fluxes plotted of particles travelling antiparallel to the field.







**Fig. 6.** Similar to Fig. 5, but for the subsequent pass over the Southern Hemisphere.

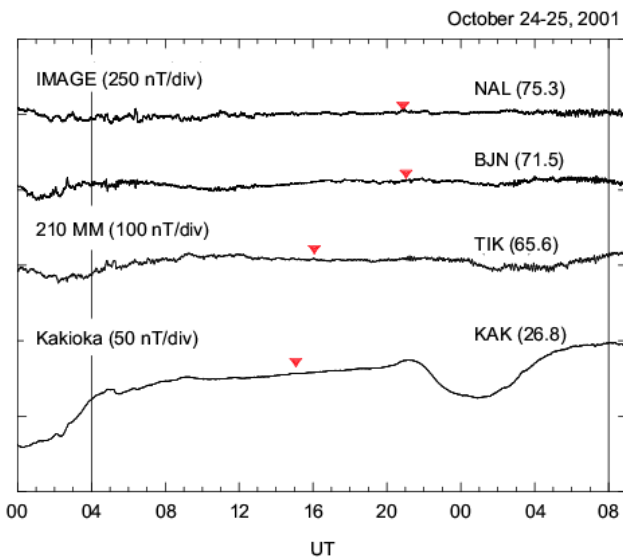
Hemispheres, respectively. In each figure the panels show from top to bottom: ion density, magnetic field deflections, electron and ion differential energy fluxes, horizontal (violet) and vertical flows. Mantle precipitation bounds a polar cap area at the dawn and dusk terminator regions. The polar rain precipitation appears tenuous and seems confined to high MLATs on the dusk side and lower MLATs on the dawnside. The electron precipitation in the polar caps occurs in energies up to  $\sim 500$  eV.

In the pass shown in Fig. 5, during which F13 climbed to  $\sim 86^\circ$ , there is evidence of a some antisunward flow confined to the dawn side of the polar cap, with little or none in the center of the polar cap. The associated potential drop is very small and estimated as  $\sim 10$  kV. Flows are less than  $300 \text{ m s}^{-1}$  except for three brief excursions when they reached  $\sim 0.5 \text{ km s}^{-1}$ . In the south (Fig. 6), where again the spacecraft reached MLATs of  $\sim 85^\circ$ , the precipitation intensity and energy range are comparable. There is evidence of a weak antisunward flow ( $\sim 200\text{--}300 \text{ m s}^{-1}$ ) at midafternoon MLTs which turns to weakly sunward outside the polar cap, thus forming a weak 2-cell DC flow as indicated also by SuperDARN measurements detailed in Sect. 2.5 below. Across this 2-cell pattern the potential drop is  $\sim 15$  kV and peak flows reach  $450 \text{ m s}^{-1}$ . Considering the region of antisunward flows, even if intermittent, as delineating the ex-

tent of the open field line region, we have in Fig. 6 the polar cap boundary located approximately at  $-80^\circ$  (dusk side) and  $-84^\circ$  (dawn side). Weak polar arcs associated with correspondingly weak flows are observed on the dusk side in the north ( $\sim 84^\circ$  MLAT) and on the dawn side in the south ( $-82^\circ$  MLAT). This arc location is in both hemispheres on the opposite sides of the moderately antisunward convection.

#### 2.4 Viscous interactions

We next investigate the possibility of viscous interaction. In statistical studies of the polar cap potential, a residual potential difference not due to reconnection is often ascribed to viscous-type processes, which are believed to contribute some 25–30 kV (e.g. Reiff et al., 1984; Cowley, 1986) or even less (Milan, 2004). This is approximately consistent with what we obtain from SuperDARN, the Boyle et al. (1997) empirical equation (Fig. 2, last panel), and the DMSP F13 passes we surveyed. However, the Kelvin-Helmholtz (KH) instability, a prime candidate thought to mediate these interactions, is apparently not operating during our interval. This is suggested by the ground magnetometer records. According to the field line resonance theory of the excitation of ULF geomagnetic pulsations by KH waves at the magnetopause and its boundary layers (Southwood, 1974; Chen and Hasegawa, 1974), this instability would give



**Fig. 7.** Stack plots of the horizontal component of the Earth's field at 4 stations. (Note that the vertical scale changes as indicated on the left). The top two stations form part of the IMAGE chain, and the bottom stations form part of the 210 MM chain. The red arrows indicate local midnight.

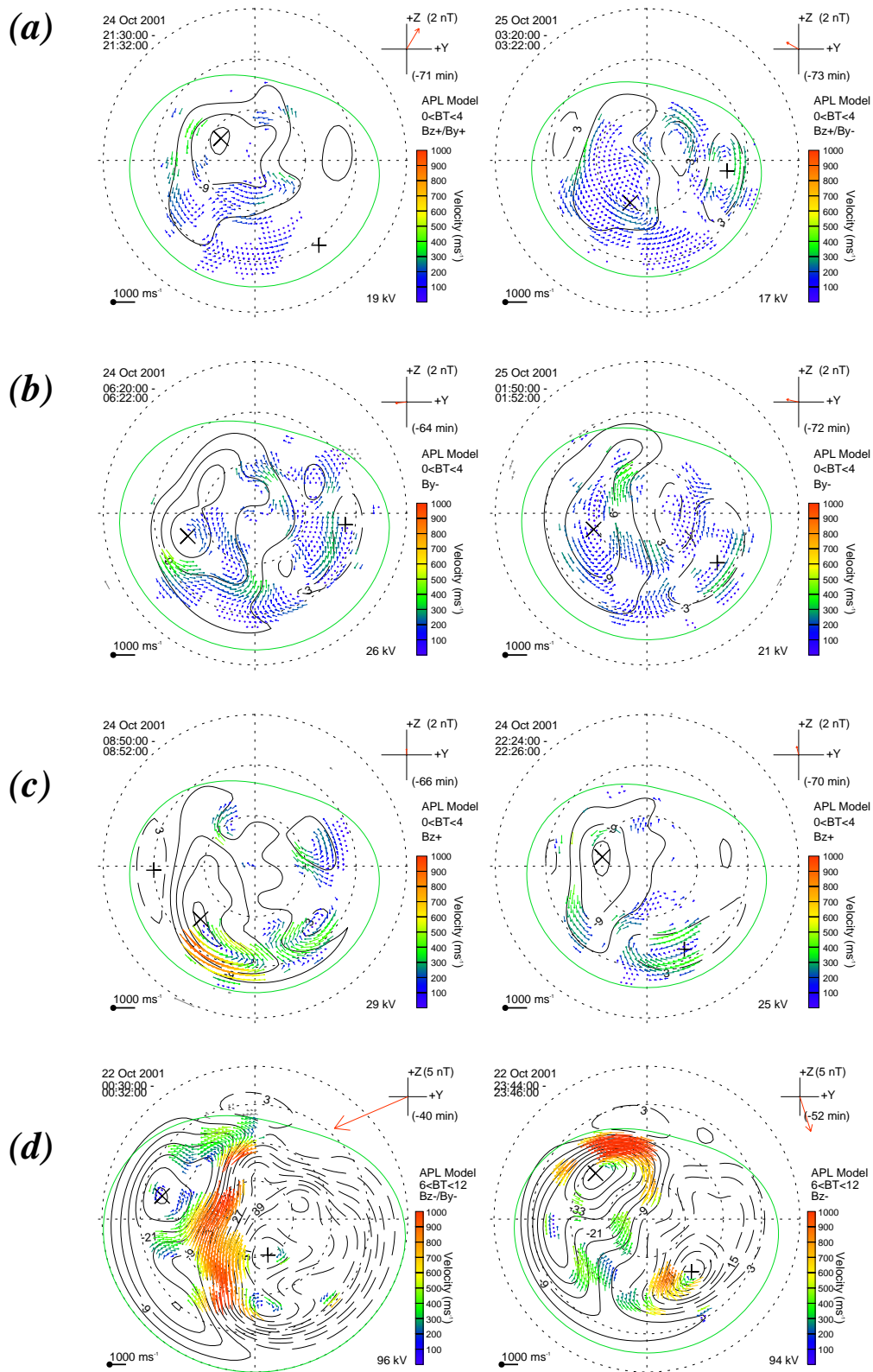
rise to geomagnetic field pulsations in the Pc 5 range (2–7 mHz). However, all ground magnetometer chains show just a quiet field. As an example, Fig. 7 displays readings of the horizontal component of the geomagnetic field at four stations arranged in order of decreasing geomagnetic latitude (marked near the station designation). The top two stations are from the IMAGE chain (Viljanen et al., 1997), and the bottom two are from the 210 deg MM array (Yumoto et al., 1992). The inverted arrows mark local magnetic midnight at the respective stations. We see neither negative bays (i.e., localized depressions of the geomagnetic field propagating poleward) at the auroral zone stations NAL and BJN, precluding the possibility of substorms, nor pulsations at the midlatitude station TIK, precluding in turn the possibility of KH instability (the irregular oscillations during early October 24 being presumably of solar wind dynamic pressure origin). The steady increase in the readings at the low-latitude station KAK reflects the steady recovery of the ring current, also indicated in the  $D_{st}^*$  profile in Fig. 2. This increase is interrupted by a wide depression centred near local noon. This is a quiet-time phenomenon (Sq-variation), thought to be the effect of neutral wind motions in the ionosphere (Nishida, 1978). (It is subtracted when the  $D_{st}$  index is compiled.) We should note, however, a possible restriction on this interpretation. Hall currents in the magnetosphere we are discussing may be weak so that the lack of geomagnetic deflections on the ground could conceivably be due to this rather than the absence of KH instability.

## 2.5 SuperDARN observations of the ionospheric flow

The Super Dual Auroral Radar Network (SuperDARN) (Greenwald et al., 1995) provides line-of-sight velocity measurements of the high-latitude ionosphere in the Northern Hemisphere. These measurements are routinely combined using the “Map Potential” fitting technique (Ruohoniemi and Baker, 1998) to produce maps of large-scale ionospheric convection. We have examined the convection patterns for the current interval and present a selection of them in Fig. 8. Each map is plotted in a magnetic latitude – magnetic local time coordinate system with noon to the top and dusk to the left. Contours of constant electric potential are shown in black, with vectors representing the fitted radar measurements shown in colour, scaled according to the bar on the right of each map. The concurrent IMF vector is shown in the top right for reference and the total transpolar voltage is given in the bottom right.

In Fig. 8a we show two maps which are representative of the nature of the flow over much of the interval. It is immediately evident that these flows do not fit within the traditional framework of convection classified in terms of the upstream IMF (as discussed by, for example, Cowley and Lockwood, 1992). This is likely to be a direct consequence of the fact that the IMF itself is not easily characterized, often having  $B_y$  and  $B_z$  components which are small and variable, as noted in Sect. 2.1. Instead, the flows are characterized by low-velocity, multi-celled patterns and regions of complete stagnation. Often, however, there is some evidence of a very low level of driving, illustrated by the moderate twin-vortex convection evidenced in Fig. 8b. There is also some enhancement to the “return”-type flows, although often the polar cap itself is essentially stationary. The transpolar voltage remains below 30 kV for much of the interval, often dipping below 20 kV (as in Fig. 8a), in general agreement with an extremely low level of dayside driving. Although data coverage is not complete, these values are also in good agreement with the Boyle empirical formula (Boyle et al., 1997; see Fig. 2 for a comparison), which, for low polar cap potentials has been repeatedly found to give a good estimate of the polar cap potential. They are also consistent with DMSP values. Between ~14:00 and 20:00 UT data coverage from the radars was particularly reduced. However, there is nothing in the supporting datasets to suggest that this period exhibited significantly different characteristics to those before and after. Overall, data coverage is remarkably good for such an extended quiet interval.

The convection maps shown in Figs. 8a–b give a good indication of the general nature of the flows throughout the interval. There are, however, a number of short lived (20–30 min) departures from this state, two examples of which are shown in Fig. 8c. These generally take the form of enhancements to the nightside auroral zone flow and resemble the flow bursts discussed by Grocott et al. (2003, 2004, 2005). Their observations were made during



**Fig. 8.** Large-scale convection inferred from the SuperDARN radars. The coordinate system is magnetic latitude versus magnetic local time. Noon is at the top and dusk is to the left. Equipotential contours are shown by the black traces. The colored vectors show fitted velocities. The IMF direction is shown at top right, and the cross-cap potential is shown at bottom right. Panels (a–c) refer to the interval we study; while panel (d) is from the earlier interval and is included for comparison.

IMF-northward, non-substorm intervals and were interpreted as signatures of reconnection in a “quiet time” tail. Although the IMF  $B_y$  and  $B_z$  components were weaker during this interval than during those studied by Grocott et al., the bursty nature of the nightside activity does imply that some low-level dayside driving is likely to have been present. It is also possible that the preceding interval of intense activity (whose associated flows are illustrated in Fig. 8d for comparison with the present observations) could well have left the magnetosphere suitably primed to unload itself in this fashion.

We now take a critical look at the use of the Map Potential technique for flows as weak as we generally have in this period. We note first that the coverage of radar data is actually quite good. Second, the actual patterns suggested by the technique may not be totally reliable, but certainly the technique is not making up the fact that the flows are very small. Third, it is the fact that the patterns are somewhat erratic – suggesting a lack of coherent and continued twin-vortex convection – which interests us here and not so much the detailed shape of the patterns themselves.

To probe our results further, we made an additional check where we reproduced the SuperDARN flow maps using a false IMF input (and hence a different statistical model input) with  $B_z = -4$  nT and found that the potential patterns changed remarkably little. Consider the right-hand map of Fig. 8a, for example. The transpolar voltage showed just a 4 kV increase (21 kV as opposed to 17 kV) with the false IMF model applied, whereas the model voltage itself for  $B_z = -4$  nT is 56 kV. For comparison, if the model were dictating the actual map shown in Fig. 8a the voltage would be 27 kV, even for northward IMF. We therefore conclude that the measured flow data is dominating the patterns and is clearly indicating convection which is weaker than might “usually” be expected.

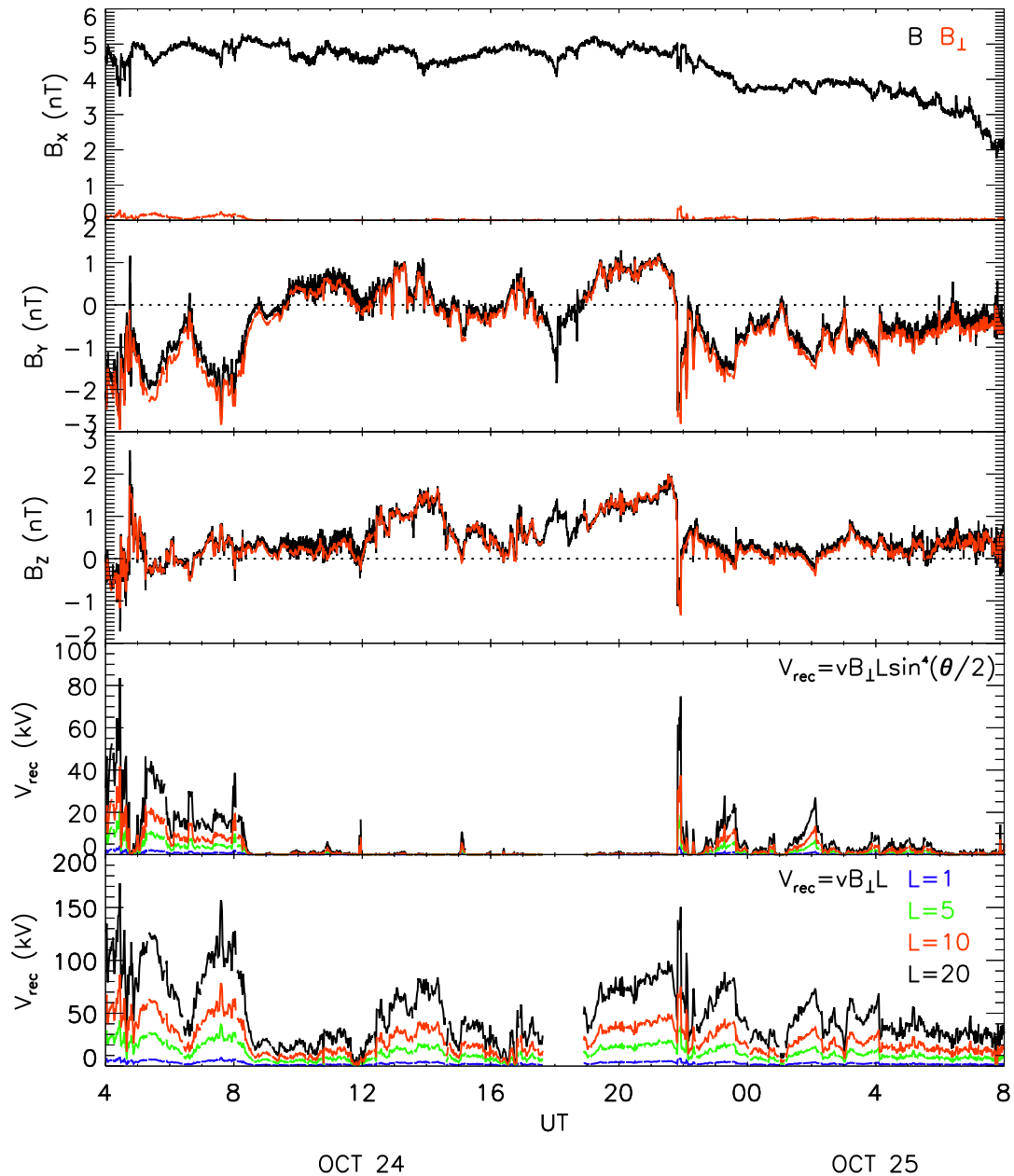
### 3 Discussion

In this paper we have discussed a most peculiar state of the magnetosphere in which the polar cap contained a quantity of open flux yet only weak, and often sporadic, convection occurred. The initial state of the magnetosphere, i.e. one which contained a moderate-to-large amount of open flux, can be easily explained by considering the previous activity caused by the passage of the ejecta merger. The leading part of the configuration, lasting for about 60 h, disturbed geospace extensively and, in particular, gave rise to two intense ( $D_{st} < -100$  nT) storms and the consequent growth of the polar cap. In contrast, the interval we have been examining elicited only weak particle precipitation in the polar cap, patchy and sometimes stagnant ionospheric convection, a monotonic decrease in the total ring current energy, as measured by the  $D_{st}$  index, and an absence of substorms. We may pictorially describe this magnetosphere as one approaching a ground state.

Let us consider the substorm recurrence rate. According to Borovsky et al. (1993), who undertook a broad survey of this quantity, the mean and standard deviation of the substorm recurrence time are  $4 \pm 2$  h. A 28-h substorm-free interval is thus several standard deviations away from the mean of the distribution. Another interesting point pertains to ring current recovery. Without competing mechanisms and in the absence of fresh injection, the recovery seen here characterizes an intrinsic property of the ring current, namely its (unforced) decay time, a quantity important in ring current models. We obtain for this a value of  $\sim 1.1$  nT h<sup>-1</sup>.

The steady nature of the polar cap showing no clear and systematic tendency to shrink or expand over such a long period also relates to a lack of solar wind driving. On the one hand, there is no substantial evidence of significant and continued flux removal from the dayside, offsetting any tendency for the polar cap to expand, and on the other, no return of flux to the dayside via substorms was monitored, offsetting in turn any tendency for it to contract. Hence our suggestion that the size of the polar cap is a remnant of activity during the earlier strong phase of ejecta passage. The polar cap size is then maintained because in the later phase, the usual processes which shrink the polar cap are generally absent. So the question is, if there is no evidence of typical dayside driving, nor of substorm activity, what is generating the low level of convection evidenced in the cross polar cap potentials shown in Fig. 2?

In their discussion of the excitation and decay of convection in the magnetosphere-ionosphere system, Cowley and Lockwood (1992) suggested that in the absence of both dayside and nightside reconnection, the flow in the system would cease irrespective of the amount of open flux present. In the present case, we have a system in which there is no clear driving mechanism, yet convection is near-continuously driven, albeit at a remarkably low rate. To investigate the possible nature of this dayside driving, we have computed the dayside reconnection voltage,  $V_{\text{rec}} = v B_{\perp} L f(\theta)$ , where  $v$  is the solar wind speed,  $B_{\perp}$  is the magnetic field perpendicular to the flow direction,  $L$  is the length of the magnetopause reconnection X-line, and  $\theta$  is the IMF clock angle (i.e. polar angle in the GSM YZ-plane), for different coupling functions  $f(\theta)$  and values of  $L$ , and these are presented in Fig. 9. The top three panels show the upstream IMF for our interval, with their solar wind-perpendicular components superimposed in red. In the bottom two panels the dayside reconnection voltages are shown for  $f(\theta) = \sin^4(\theta/2)$  and  $f(\theta) = 1$ , for values of  $L$  set at 1, 5, 10 and 20  $R_E$ . Firstly, it is clear from the values of  $B_{\perp x}$  that the magnetic field was essentially parallel to the solar wind flow throughout the interval. Using this to compute the reconnection voltage (panels 4 and 5) then gives an indication of the possible coupling between the solar wind and magnetosphere. Looking first at panel 4, it is clear that conventional approximations of the dayside reconnection voltage imply essentially zero coupling for many hours over the interval. If this were true, however, then the  $\sim 20$  kV



**Fig. 9.** The top three panels reproduce the components of the IMF, on top of which are shown in red the components of  $\mathbf{B}$  perpendicular to the flow vector. The bottom two panels show two theoretical estimates of the reconnection rate parametrized by the length of the magnetopause reconnection X-line,  $L$

cross polar cap voltage evident in the radar data would, in the absence of tail reconnection, have to be due entirely to viscous coupling. As was shown in Fig. 8, there is often some evidence of convection across the polar cap, which cannot be due to a viscous interaction since this occurs only on closed field lines. If the convection was due to tail reconnection then a voltage of  $\sim 20$  kV would easily close even a large polar cap over an interval as long as we are dealing with here. The fact that this does not occur provides further evidence that

some ongoing quasi-balanced dayside and nightside convective driving must be occurring.

Panel 5 of Fig. 9 shows the theoretical maximum reconnection voltage for various  $L$  values if we were to assume that the total solar wind electric field could couple to the magnetosphere irrespective of the orientation of  $B_{\perp}$ . This shows that, for  $L=5-10 R_E$ , voltages similar to the cross polar cap potentials discussed above could occur. We can only assume therefore that some dayside reconnection geometry

with which we are not familiar is facilitating a low level of coupling in this case. Combining panels (4) and (5) leads us to suggest that, under these special circumstances,  $L$  (the geoeffective length) is large (of order  $\sim 10$ ) and the dependence on clock angle,  $\theta$  (i.e., the coupling function) is less strong than the  $\sin^4(\theta/2)$ , the Perreault-Akasofu (1978) function often employed. Then we would be able to recover the  $\sim 20$  kV voltage that we see.

The nature of the nightside component of the convection is more readily identifiable, and resembles the Grocott et al. (2003, 2004, 2005) nightside non-substorm reconnection bursts mentioned above. This form of tail reconnection was initiated by the continued driving of the solar wind under the influence of a strong and constant IMF  $B_y$  component, which is absent in the present case. However, a tendency of the tail to respond to a modest driving of any kind in this fashion may be responsible for the present observations. In summary, the trailing end of this complex ejecta may be characterized by the general absence of a coherent driver which is reflected in the lack of coherent and continued twin-vortical ionospheric convection, the absence of substorms, a background polar cap potential and weak and occasionally absent polar cap precipitation.

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